

On the evolution of the resonant planetary system HD 128311

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ABSTRACT

Context. A significant number of the known multiple exoplanetary systems are containing a pair of giant planets engaged in a low order mean motion resonance. Such a resonant condition protects the dynamics of these planets resulting in very stable orbits. According to recent studies the capture into a resonance is the result of a planetary migration process induced by the interaction of the planets with a protoplanetary disk. If the migration is slow enough (adiabatic) next to a mean motion resonance, the two planets will also be in apsidal corotation.

Aims. The recently refined orbital parameters of the system HD 128311 suggest that the two giant planets are in a 2:1 mean motion resonance, however without exhibiting apsidal corotation. Thus the evolution of this system can not be described by an adiabatic migration process alone. We present possible evolution scenarios of this system combining migration processes and sudden perturbations.

Methods. We model migration scenarios through numerical integration of the gravitational N-body problem with additional non-conservative forces. Planet-planet scattering has been investigated by N-body simulations.

Results. We show that the present dynamical state of the system HD 128311 may be explained by such evolutionary processes.

Key words. planets and satellites: formation - celestial mechanics - methods: N-body simulations

1. Introduction

Among the 19 multi-planet systems found to date about a third are engaged in a low order mean motion resonance (Vogt et al. 2005). The most prominent case is the exact 2:1 resonance of the two outer planets in GJ 876. There the orbital elements are very well determined, due to the short periods of the planets of ≈ 30 and 60 days (Laughlin et al. 2005). The formation of resonant configurations between planets must be due to dissipative processes altering the semi-major axis. For the system GJ 876 the orbits are in apsidal corotation and both resonant angles librate with small amplitudes, a condition which can best explained by a sufficiently slow and long lasting differential migration process induced by the interaction of the planets with a protoplanetary disk (Snellgrove et al. 2001; Lee & Peale 2002; Kley et al. 2005). Hence, the occurrence of resonances constitutes a very strong indication of migration in young planetary systems, in addition to the hot Jupiter cases.

Recent analysis of the planetary system HD 128311 suggests that two giant planets are engulfed in a 2:1 mean motion resonance, however without exhibiting apsidal corotation (Vogt et al. 2005). On the other hand, according to Lee & Peale

(2002); Beaugé et al. (2006), in the case of a sufficiently slow migration process the resonant planets should also be in apsidal corotation. In order to resolve the above discrepancy we construct *mixed* evolutionary scenarios of migrating planetary systems incorporating migration and other additional perturbative effects. For our investigation we use N-body numerical integrations containing also non-conservative forces (Lee & Peale 2002), which have been tested on full hydrodynamic evolutions of embedded planets (Kley et al. 2004). In this letter we report our findings in modeling the behavior of the resonant exoplanetary system HD 128311.

2. Orbital solution and its stability

It is mentioned by Vogt et al. (2005) that the best fit to the radial velocity curve of HD 128311, based only on the use of Keplerian ellipses, leads to orbital data resulting in unstable behavior. Therefore they present also an alternate fit for HD 128311 calculated by using three-body gravitational interactions. The orbital data determined by this alternate fit (listed in Table 1) result in stable orbits, where the two giants are in a protecting 2:1 mean motion resonance (MMR). We perform three-body numerical integrations using these initial conditions

Table 1. Orbital data of HD 128311 as provided by Vogt et al. (2005). M denotes the mean anomaly.

planet	mass [M_J]	a [AU]	e	M [deg]	ϖ [deg]
b	1.56	1.109	0.38	257.6	80.1
c	3.08	1.735	0.21	166.0	21.6

of Vogt et al. (2005), as listed in Table 1. Only the resonant angle $\Theta_1 = 2\lambda_2 - \lambda_1 - \varpi_1$ librates around 0° with an amplitude of $\sim 60^\circ$ (Fig. 1, top), while both $\Theta_2 = 2\lambda_2 - \lambda_1 - \varpi_2$ and $\Delta\varpi = \varpi_2 - \varpi_1$ circulate. Here λ_j are mean longitudes and ϖ_j are longitudes of periaapse, both numbered from the inside out. This implies that the systems, although engaged in a MMR, is not in apsidal corotation. Moreover, the eccentricities show also large oscillations (Fig. 1, bottom).

As described by Vogt et al. (2005), due to the relatively high stellar jitter, there exist a large variety of dynamically distinct stable orbital solutions resulting in the current radial velocity curve. The authors present the results obtained by a self-consistent two-planet model (taking into account the mutual gravitational interactions) and a Monte-Carlo procedure. As final outcome, they have found that only very few samples exhibit stable behavior without resonance. The remaining stable solutions are in 2:1 MMR, in the majority of the cases ($\sim 60\%$) only Θ_1 librates while Θ_2 circulates, in the remaining cases Θ_2 librates as well. Hence, we assume that their derived parameters in Table 1 are an accurate representation of the dynamics of the system.

3. Evolution through a migration process

3.1. Adiabatic migration

As shown in the previous section, the planets in the system HD 128311 are presently engaged in a stable 2:1 MMR. Although the actual orbital parameters do not exhibit apsidal corotation, the existence of the resonance suggests that the system has in the past gone through a migration process.

The migration of a single planet can be characterized by the migration rate \dot{a}/a and the eccentricity damping rate \dot{e}/e . Here, we use the corresponding e-folding times for the semi-major axes and eccentricities: τ_a and τ_e , respectively. The relation between the damping rates and e-folding times is $\dot{a}/a = -1/\tau_a$, and similarly for the eccentricities. Investigating the system GJ 876, Lee & Peale (2002) have found that for a sufficiently slow migration the final state of the system depends only on the ratio K of the e-folding times $K = \tau_a/\tau_e$.

From hydrodynamical calculations we know that the order of magnitude of K is close to unity and it reflects the physical properties (i.e. mass and viscosity) of the protoplanetary disk (Kley et al. 2004). If a system is a subject to such an adiabatic migration, a given value of K results in unique values of the final eccentricities (Lee & Peale 2002). However, as shown above, in the case of HD 128311 the eccentricities are varying with quite large amplitudes.

We have performed a set of numerical integrations of the general co-planar three-body problem adding non-conservative drag forces, varying the value of K .

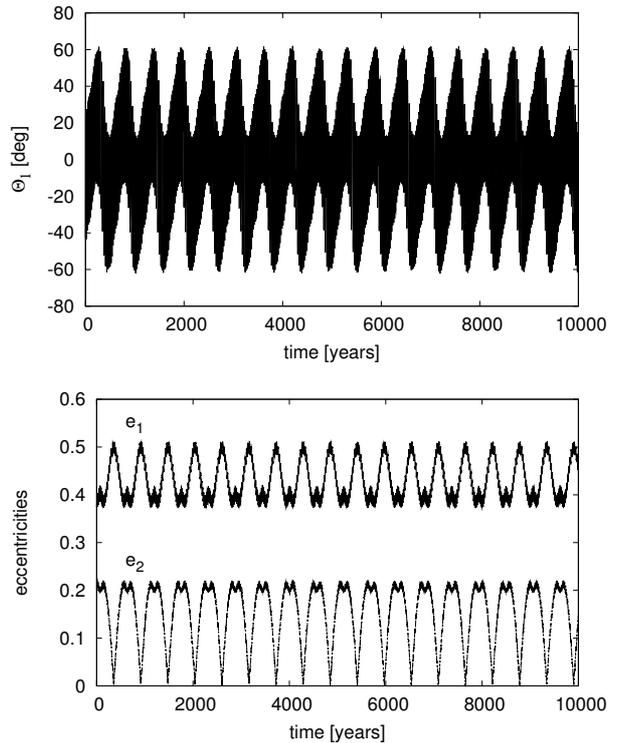


Fig. 1. Top: Time evolution of the resonant angle Θ_1 obtained by numerical integration using initial conditions of Table 1. **Bottom:** The corresponding evolution of the eccentricities.

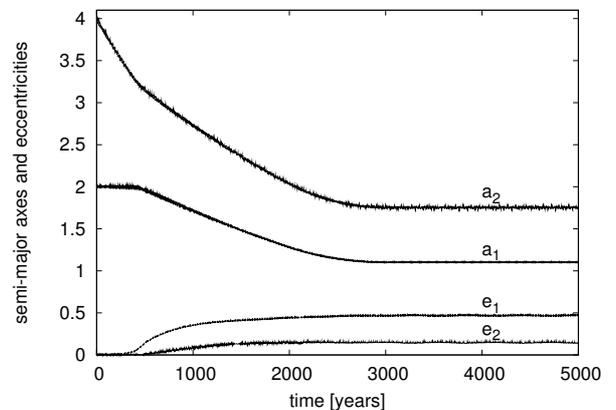


Fig. 2. The behavior of the semi-major axes and eccentricities during an adiabatic migration with $\tau_a = 2 \times 10^3$ years and $K = 5$. The migration is stopped between 2×10^3 and 3×10^3 years by applying a linear reduction.

In these simulations the planets move originally in quasi-circular orbits and we start them from $a_1 = 4\text{AU}$ and $a_2 = 2\text{AU}$. Assuming that only the outer planet is forced to migrate inward, we implement a dissipative force which results in a migration characterized by the e-folding time $\tau_a = 2 \times 10^3$ years in the semi-major axis, and $\tau_e = 4 \times 10^2$ years in the eccentricity of the outer planet (corresponding to $K = 5$). In order to obtain the present values of the semi-major axes we slow down the migration beginning from $t_1 = 2 \times 10^3$ years and decreasing it linearly to zero until $t_2 = 3 \times 10^3$ years. In this way we model the smooth dispersal of the protoplanetary disk. After

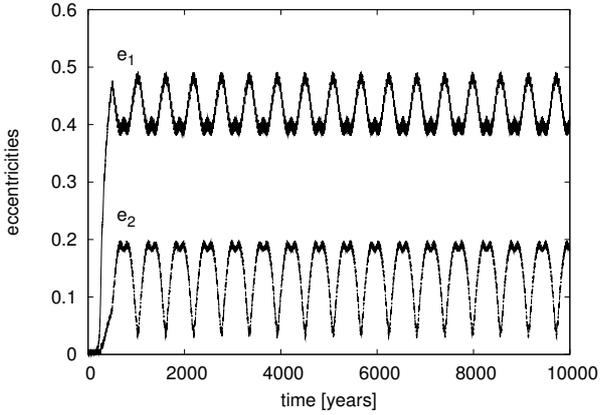


Fig. 3. The behavior of the eccentricities obtained by a sudden stop of the migration of the outer planet.

this migration process, the system is locked into a deep 2:1 resonance, the resonant angles $\Theta_{1,2}$, and $\Delta\varpi$ librate around 0° with small amplitudes. The eccentricities e_1 and e_2 are almost constant (showing only a small amplitude oscillations), where the index 1 refers to the inner and the index 2 to the outer planet. For the chosen $K = 5$ we find $e_1 = 0.46$ and $e_2 = 0.15$, see Fig. 2. Smaller/larger values of K yield always systems deep in resonance and result in larger/smaller $e_{1,2}$, contradicting Fig. 1.

Clearly, the present behavior of HD 128311 is not the result of such an adiabatic migration process alone. In the following, we shall present two additional mechanisms, which may be responsible for the large oscillations of the eccentricities and breaking the apsidal corotation.

3.2. Sudden stop of the migration of the outer planet

Recent *Spitzer* observations of young stars confirm that the inner part of the protoplanetary disk may contain only very little mass D'Alessio et al. (2005); Calvet et al. (2005), possibly due to photo-evaporation induced by the central star. Thus, upon approaching the inner rim of such a disk, the inward migration of a planet can be terminated rapidly.

In order to model this type of scenario, we perform additional simulations where the migration of the outer planet has been stopped abruptly reaching the actual value of its semi-major axis ($a_2 = 1.73\text{AU}$). We assume that the inner planet orbits in the empty region of the disk at $a_1 = 1.5\text{AU}$, and we start the outer planet from $a_2 = 4\text{AU}$ forcing to migrate inward very fast ($\tau_a = 500$ years). We find that the present behavior of the eccentricities can be obtained by using $K = 10$. The sudden stop of the migration results in a behavior of the eccentricities shown in Fig. 3 which is very similar to the observed case (Fig. 1). We note that after the above sudden stop of the migration the planets remain in apsidal corotation but with increased libration amplitudes of the resonant angles.

3.3. Planet-planet scattering event

The behavior of the eccentricities of the giant planets of the system HD 128311 is very similar to that observed in the system

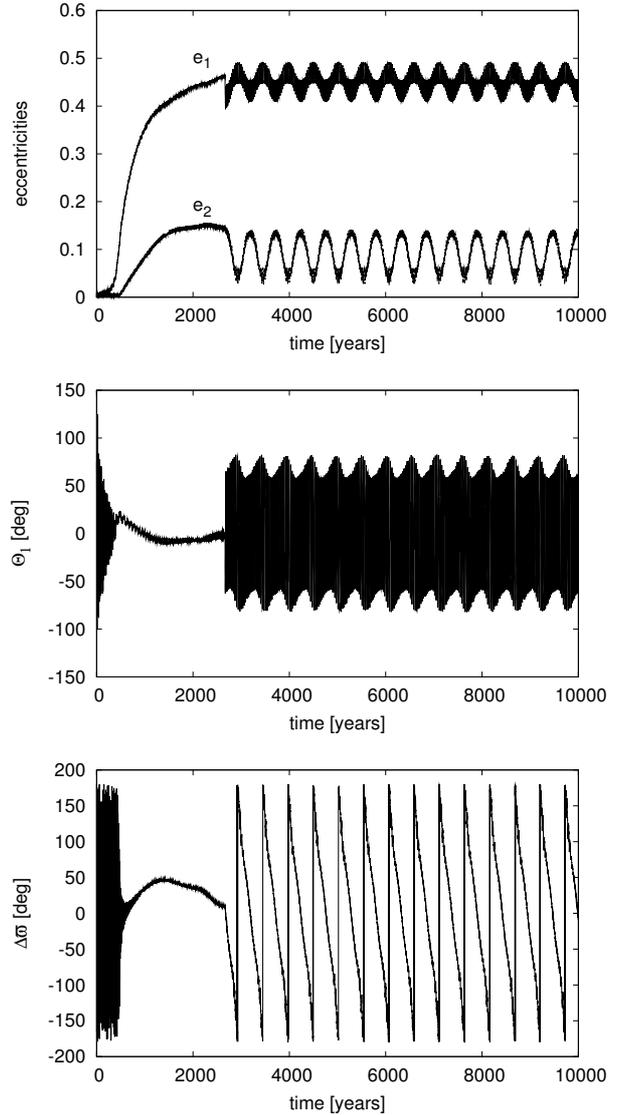


Fig. 4. Time evolution of the eccentricities (**Top**), resonant angles Θ_1 (**Middle**) and $\Delta\varpi$ (**Bottom**) of the inward migrating giant planets, before and after a scattering event with an inner low mass planet.

around ν Andromedae. Ford et al. (2005) proposed that such a behavior in ν And is most likely the result of a planet-planet scattering event. Investigating the resonant system HD 73526, Tinney et al. (2006) also suggested that the lack of the apsidal corotation can be the result of a dynamical scattering event. Therefore we investigate whether the present behavior of HD 128311 can be modeled by such an effect. We present two cases.

Additional inner planet

In this case we assume that an additional small mass planet is orbiting close to the central star in a quasi circular orbit. When the outer giants are far enough from the star the orbit of the small mass planet is stable. However, as the giants migrate inward approaching their present positions they perturb the orbit of the small planet making its motion chaotic, which in long term may result in an increase of its eccentricity. Finally, due

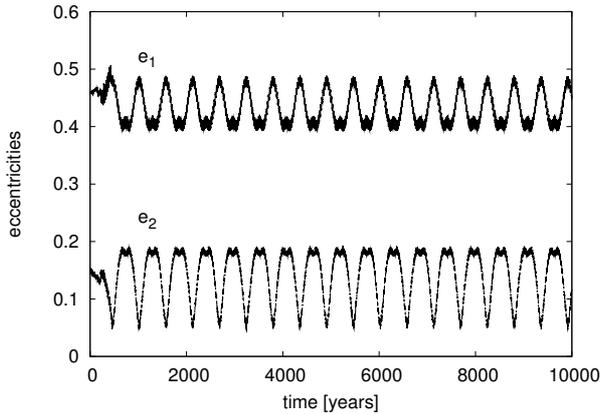


Fig. 5. The behavior of the eccentricities before and after a scattering with a small mass planet migrating inward from outside starting from $a = 2.6\text{AU}$.

to its high eccentricity, the small planet can suffer a close encounter with one of the giant planets. After the close encounter the small planet may be ejected from the system or pushed into an orbit with very large semi-major axis. Depending on the initial position of the small planet, it can also be captured into a mean-motion resonance with the inner giant, and forced to migrate inward. During the migration its eccentricity will increase in a much shorter timescale than by a “pure” chaotic evolution. Thus, in this case a close encounter with the inner giant is very likely as well.

In Fig. 4 we show the behavior of the eccentricities and the resonant angles Θ_1 and $\Delta\varpi$ when the giant planets are migrating inwards with initial conditions given in Sect. 3.1. Starting the small planet ($m = 0.03M_{\text{Jup}}$) from a nearly circular orbit with $a = 0.5\text{AU}$, it is captured into a resonance with the inner giant. During the migration the small planet’s eccentricity increases, and finally the small planet crosses the orbits of giant planets and suffers a close encounter with them. After the encounter at $t \approx 2500\text{yrs}$, the apsidal corotation of the giant planets breaks (i.e. Θ_2 and $\Delta\varpi$ circulate), however the giant planets remain in the 2:1 resonance, since Θ_1 still librates around 0° , see Fig. 4.

Additional outer planet

In this second case we assume that the third planet ($m = 0.03M_{\text{Jup}}$) originates from the outer region, and that it approaches the giant planets through inward migration where we assume that the adiabatic migration of the two giant planets is already finished, as displayed in Fig. 2. They orbit for instance in a gas-free environment while the small planet is still embedded in the outer protoplanetary disk. In Fig. 5 we display the variation of the eccentricities of the giants after the scattering event. The small planet is started from $a = 2.6\text{AU}$ and migrates inward with an e-folding time $\tau_a = 2 \times 10^3$ years. After scattering it is pushed into a very distant orbit $a \sim 300\text{AU}$. We note that in this case the giant planets remain in apsidal corotation, however with a substantial increase in Θ_2 and $\Delta\varpi$. In additional simulations we find that apsidal resonance is always preserved, however we cannot exclude the possibility of breaking the apsidal corotation entirely.

4. Conclusions

In this paper we investigate evolutionary scenarios, which may have led to the observed behavior of the resonant system HD 128311, where the two planets are engaged in a 2:1 mean motion resonance, but not in apsidal corotation. We assume that the two giant planets have been formed far from the central star and migrated inwards, due to gravitational interaction with the protoplanetary disk. During this differential migration, they have been locked into the 2:1 mean motion resonance. At the end of their migration we imagine that the system suffers a sudden perturbation.

As a first case we study the sudden stop of the migration, which may possibly be induced by an inner rim of the disk and an empty region inside of it, as indicated by some observations of young stars.

In the second scenario the sudden perturbation is caused by a planet-planet scattering event, similar to that suggested by Ford et al. (2005) in the case of ν And. We analyze an encounter with a small ($\sim 10M_\oplus$) planet, approaching the two massive planets either from outside or inside, which is thrown to a large a orbit or directly ejected after the scattering event.

We find that the sudden perturbation caused by an encounter with an *inner* small planet is clearly able to break the apsidal resonance of the two planets. But also in the other cases the orbital behavior of the giants is very similar to the most recent orbital data of HD 128311 (Vogt et al. 2005).

The system HD 128311 constitutes another example which demonstrates the important interplay of migration and scattering processes in shaping the dynamics of exoplanetary systems.

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